Supernova neutrinos

Collective and matter effects at large detectors

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Changing paradigm of SN neutrino oscillations

Neutrino-electron forward scattering: MSW effects (1999 –)

- Flavour conversions mainly in MSW resonance regions : ($\rho \sim 10^{3-4}$ g/cc, 1–10 g/cc)
- Sensitivity to $\sin^2\theta_{13} \gtrsim 10^{-5}$ and mass hierarchy

Neutrino-neutrino forward scattering: Collective effects (2006 –)

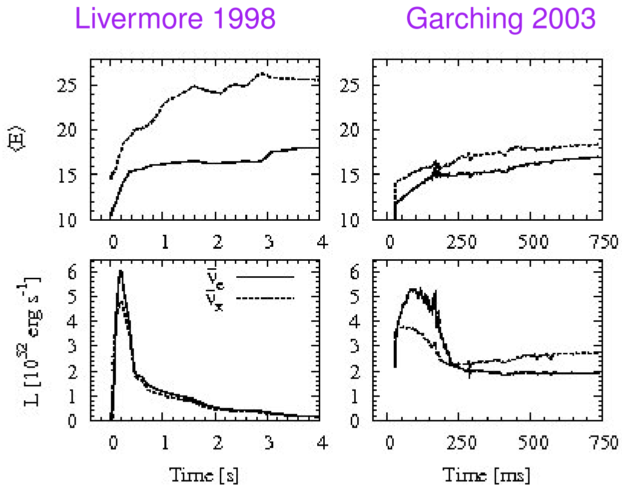
- Significant flavour conversions near the neutrinosphere : ($\rho \sim 10^{6-10}$ g/cc)
- Synchronized osc $\rightarrow$ bipolar osc $\rightarrow$ spectral split
- Single spectral split: In IH,
  $\bar{\nu}_e$ and $\bar{\nu}_\mu$ spectra swap completely
  ν_e and ν_μ spectra swap for $E > E_c$
- Sensitivity even to $\sin^2\theta_{13} \sim 10^{-10}$

Changing paradigm of SN neutrino oscillations

Multiple spectral splits (2009 –)

- "Single spectral split" valid only when $L_{\nu_e} \approx L_{\bar{\nu}_e} \gtrsim L_{\nu_\mu}$
- In general, both $\nu_e \leftrightarrow \nu_y$ and $\bar{\nu}_e \leftrightarrow \bar{\nu}_y$ swaps take place, in sharply separated energy regions
$$\begin{pmatrix} \nu_x \\ \nu_y \end{pmatrix} = \begin{pmatrix} \cos\theta_{23} & \sin\theta_{23} \\ -\sin\theta_{23} & \cos\theta_{23} \end{pmatrix} \begin{pmatrix} \nu_\mu \\ \nu_\tau \end{pmatrix}$$
- Three flavour effects: even $\nu_e \leftrightarrow \nu_x$ and $\bar{\nu}_e \leftrightarrow \bar{\nu}_x$ swaps take place, in sharply separated energy regions
- The swapped / unswapped energy regions depend on primary fluxes and mass hierarchy

Primary neutrino fluxes: a lot of model dependence



Livermore 1998

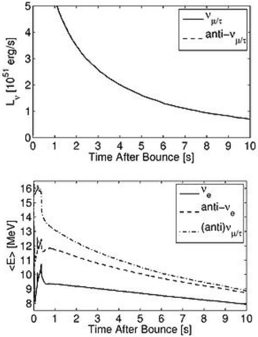
Garching 2003

solid: $\bar{\nu}_e$, dotted: $\bar{\nu}_x$

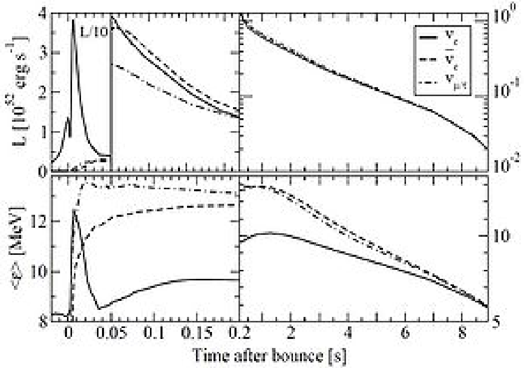
Totani et al., 1998, Raffelt et al., 2003

Recent model preditions for fluxes

Basel 2009



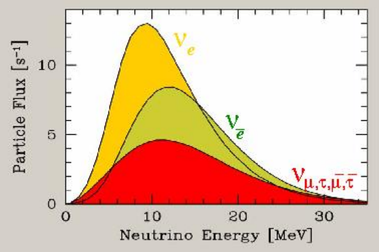
Garching 2009



Fischer et al, 2009; Hüdepohl et al, 2009

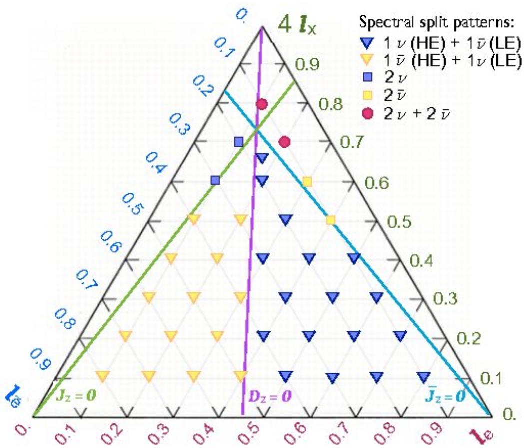
- Average energies slightly smaller
- $\sim$ 20% differences in average energies and fluxes (especially during the accretion phase; more for neutrinos)

Typical features of the spectra



- Average energies:
  $\langle E_0(\nu_e)\rangle \approx 10\text{–}12$ MeV
  $\langle E_0(\bar{\nu}_e)\rangle \approx 13\text{–}16$ MeV
  $\langle E_0(\nu_x)\rangle \approx 15\text{–}25$ MeV
- Energy hierarchy:
  $\langle E_0(\nu_e)\rangle < \langle E_0(\bar{\nu}_e)\rangle < \langle E_0(\nu_x)\rangle = \langle E_0(\nu_y)\rangle$
- Luminosities:
  $L_{\nu_e} \approx L_{\bar{\nu}_e}$
  $L_{\nu_x} \approx (0.5 - 2.0)\ L_{\nu_e}$

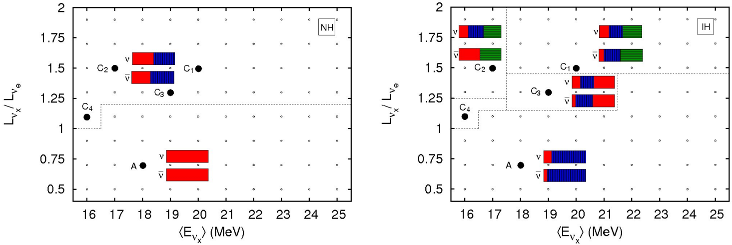
Number and position of splits: is there a pattern ?



(Inverted hierarchy, $\nu_e \leftrightarrow \nu_y$ split)

G.Fogli, E.Lisi, A.Marrone, I.Tamborra, JCAP 0910:002 (2009)

Split patterns with $\langle E_{\nu_\mu} \rangle$ and L_{ν_μ}: two “phases”



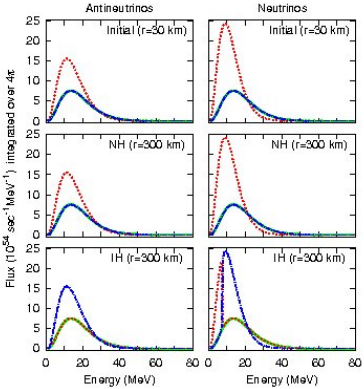
No swap, $e \leftrightarrow y$ swap, $e \leftrightarrow x$ swap

S. Choubey, B. Dasgupta, AD, A. Mirizzi, arXiv:1008.0308 [hep-ph]

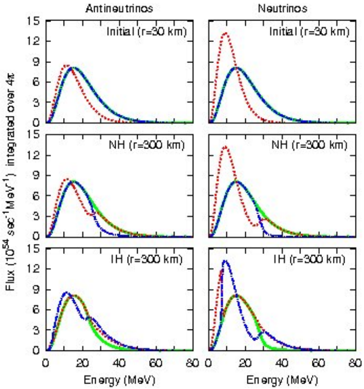
- $\langle E_{\nu_e} \rangle = 12$ MeV, $\langle E_{\bar{\nu}_e} \rangle = 15$ MeV
- $L_{\nu_e} = L_{\bar{\nu}_e}$
- For lower $\langle E_{\nu_e} \rangle$, scale $\langle E_{\nu_\mu} \rangle$ appropriately

Multiple splits in different phases

Phase A



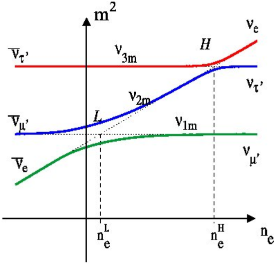
Phase C



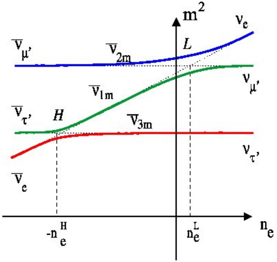
Flavours: $\nu_e, \bar{\nu}_e, \nu_y, \bar{\nu}_y, \nu_x, \bar{\nu}_x$

MSW Resonances inside a SN

Normal mass ordering



Inverted mass ordering



AD, A.Smirnov, PRD62, 033007 (2000)

H resonance: $(\Delta m^2_{\rm atm}, \theta_{13})$, $\rho \sim 10^3$–10^4 g/cc

- In $\nu(\bar{\nu})$ for normal (inverted) hierarchy
- Adiabatic (non-adiabatic) for $\sin^2\theta_{13} \gtrsim 10^{-3} (\lesssim 10^{-5})$

L resonance: $(\Delta m^2_\odot, \theta_\odot)$, $\rho \sim 10$–100 g/cc

- Always adiabatic, always in ν

Fluxes arriving at the Earth

$$F_{\nu_e} = p\, F^0_{\nu_e} + (1-p)\, F^0_{\nu_x}\,, \qquad F_{\bar{\nu}_e} = \bar{p}\, F^0_{\bar{\nu}_e} + (1-\bar{p})\, F^0_{\nu_x}$$

p at low, intermediate, high energies

		Phase A ($L_{\nu_e} \gtrsim L_{\nu_x}$)			Phase C ($L_{\nu_e} \gtrsim L_{\nu_x}$)		
NH	$\sin^2\theta_{13} \gtrsim 10^{-3}$	0	0	0	0	0	s^2
	$\sin^2\theta_{13} \lesssim 10^{-5}$	s^2	s^2	s^2	s^2	s^2	0
IH	$\sin^2\theta_{13} \gtrsim 10^{-3}$	s^2	0	0	s^2	0	c^2 (s^2)
	$\sin^2\theta_{13} \lesssim 10^{-5}$	s^2	0	0	s^2	0	c^2 (s^2)

$\bar{p}$ at low, intermediate, high energies

		Phase A ($L_{\nu_e} \gtrsim L_{\nu_x}$)			Phase C ($L_{\nu_e} \gtrsim L_{\nu_x}$)		
NH	$\sin^2\theta_{13} \gtrsim 10^{-3}$	c^2	c^2	c^2	c^2	c^2	0
	$\sin^2\theta_{13} \lesssim 10^{-5}$	c^2	c^2	c^2	c^2	c^2	0
IH	$\sin^2\theta_{13} \gtrsim 10^{-3}$	0	c^2	c^2	0	c^2 $[0]$	s^2 (0)
	$\sin^2\theta_{13} \lesssim 10^{-5}$	c^2	0	0	c^2	0 $[c^2]$	s^2 (c^2)

$s^2 \equiv \sin^2\theta_{12}, c^2 \equiv \cos^2\theta_{12}$

$(C_3, C_4), [C_2, C_4]$

Signal expected from a galactic SN (10 kpc)

Water Cherenkov detector: (events at SK)

- $\bar{\nu}_e p \rightarrow n e^+$: $\approx$ 7000 – 12000

  $$\Delta_{\mathrm{WC}}/\mathrm{MeV} = 0.47\sqrt{E_e/\mathrm{MeV}}$$

- $\nu e^- \rightarrow \nu e^-$: $\approx$ 200 – 300
- $\nu_e +^{16} O \rightarrow X + e^-$: $\approx$ 150–800

Carbon-based scintillation detector:

- $\bar{\nu}_e p \rightarrow n e^+$ ($\sim$ 300 per kt)

  $$\Delta_{\mathrm{SC}}/\mathrm{MeV} = 0.075\sqrt{E_e/\mathrm{MeV}}$$

- $\nu + {}^{12}C \rightarrow \nu + X + \gamma$ (15.11 MeV)

Liquid Argon detector:

- $\nu_e + {}^{40}Ar \rightarrow {}^{40}K^* + e^-$ ($\sim$ 300 per kt)

  $$\Delta_{\mathrm{LAr}}/\mathrm{MeV} = 0.11\sqrt{E_e/\mathrm{MeV}} + 0.02\, E_e/\mathrm{MeV}$$

Earth matter effects

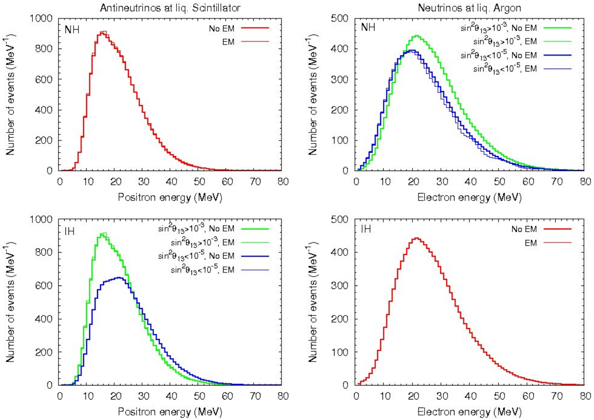
- If F_{ν_1} and F_{ν_2} reach the earth,

$$F^D_{\nu_e}(L) - F^D_{\nu_e}(0) \;=\; (F_{\nu_2} - F_{\nu_1}) \times \sin 2\theta^{\oplus}_{12} \sin(2\theta^{\oplus}_{12} - 2\theta_{12}) \sin^2\left(\frac{\Delta m^2_{\oplus} L}{4E}\right)$$

(Sign changes for antineutrinos)

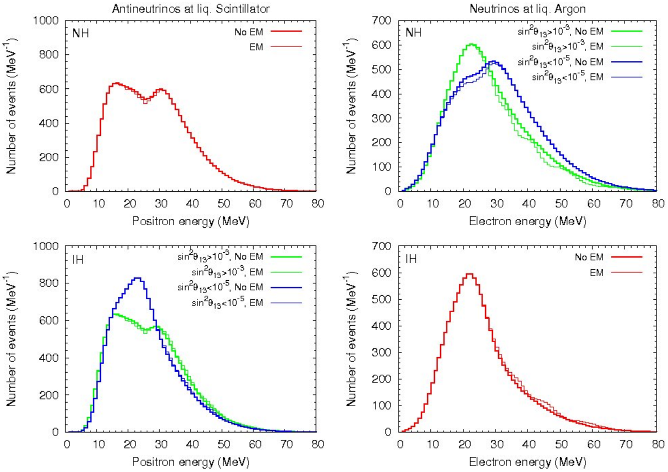
- Nonzero Earth matter effects require
  - Neutrinos: $p \neq 0$
  - Antineutrinos: $\bar{p} \neq 0$
- Can distinguish scenarios depending on Earth effects in different energy regimes
- A more efficient way of detecting split positions

Spectra at detectors with Earth effects: phase A



- Spectral splits not visible
- Earth effects possibly visible in neutrinos

Spectra at detectors with Earth effects: phase C



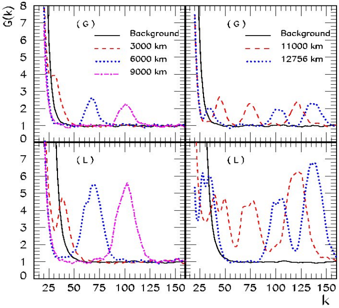
- Spectral split may be visible as “shoulders”
- Earth effects possibly visible, more prominent in ν_e

Earth effects: oscillations at a single detector

Fourier power spectrum: $G_N(k) = \frac{1}{N} \left| \sum_{events} e^{iky} \right|^2$

$(y \equiv 25\ MeV/E)$

- Peak positions model independent, at known frequencies



AD, M. Kachelrieß, G. Raffelt,

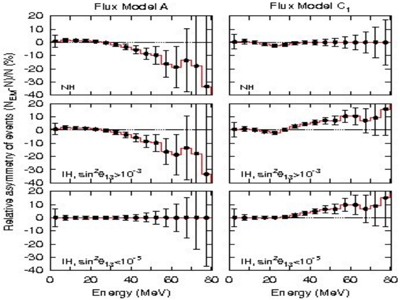
R. Tomàs, JCAP 0401:004 (2004)

Comparison between two detectors

- Ratio of luminosities at IceCube and a megaton water Cherenkov, as a function of time

AD, M. Keil, G. Raffelt, JCAP 0306:005 (2003)

- Comparing spectra at two 400 kt water Cherenkovs



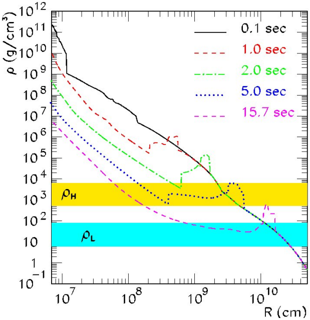
S. Choubey et al., arXiv:1008.0308 [hep-ph]

Robust experimental signature

- Earth effects can distinguish hierarchies even for $\theta_{13} \lesssim 10^{-10}$

Shock wave and adiabaticity breaking

When shock wave passes through a resonance region (density ρ_H or ρ_L):



- adiabatic resonances may become momentarily non-adiabatic
- Sharp changes in the final spectra even if the primary spectra change smoothly

R. C. Schirato, G. M. Fuller, astro-ph/0205390

G. L. Fogli, E. Lisi, D. Montanino and A. Mirizzi, PRD 68, 033005 (2003)

J.P.Kneller, G.C.Mclaughlin, J.Brockman, PRD77, 045023 (2008)

Shock signals at a megaton water Cherenkov

- Time-dependent dip/peak features in $N_{\nu_e,\bar{\nu}_e}(E)$, $\langle E_{\nu_e,\bar{\nu}_e}\rangle$, $\langle E^2_{\nu_e,\bar{\nu}_e}\rangle$, etc.
- Times at which dips/peaks appear in $N_{\bar{\nu}_e}(E)$ are the times at which the shock waves enter the densities

$$\rho(E) = \frac{m_N \Delta m^2_{atm}}{2\sqrt{2} G_F Y_e E}$$

- $\Rightarrow$ Tracking of shock wave while it is still inside the mantle

R.Tomas, M.Kachelriess, G.Raffelt, AD, H.T.Janka and L.Scheck, JCAP **0409**, 015 (2004)

Identifying mixing scenario

- Shock wave present in ν_e only for NH $\oplus \sin^2\theta_{13} \gtrsim 10^{-5}$
- Shock wave present in $\bar{\nu}_e$ only for IH $\oplus \sin^2\theta_{13} \gtrsim 10^{-5}$

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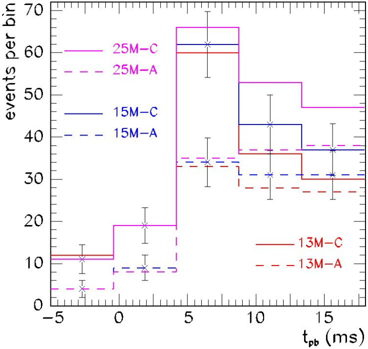
Vanishing neutronization (ν_e) burst



M. Kachelriess, R. Tomas, R. Buras,

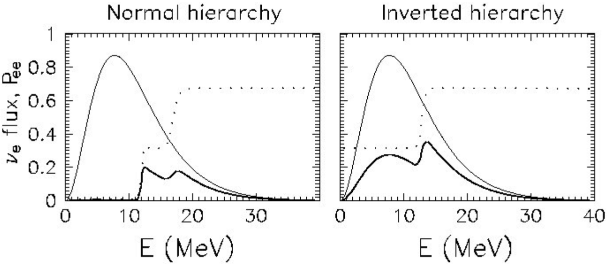
H. T. Janka, A. Marek and M. Rampp

PRD **71**, 063003 (2005)

- Time resolution of the detector crucial for separating ν_e burst from the accretion phase signal

Burst signal vanishes for NH $\oplus \sin^2\theta_{13} \gtrsim 10^{-3}$

Stepwise spectral split in O-Ne-Mg supernovae



Neutronization burst

- MSW resonances deep inside collective regions

H. Duan, G. M. Fuller, J. Carlson, Y.Z.Qian, PRL100, 021101 (2008)

C. Lunardini, B. Mueller, H. T. Janka, arXiv:0712.3000

- “MSW-prepared” spectral splits: two for IH, one for NH

H.Duan, G.Fuller, Y.Z.Qian, PRD77, 085016 (2008)

- Positions of splits fixed by initial spectra

B.Dasgupta, AD, A. Mirizzi, G.G.Raffelt, PRD77, 1130007 (2008)

- ν_e suppression more at low energy: Ar detector crucial
- Identification of O-Ne-Mg supernova ??

Inverse SN neutrino problem

Multiple independent signals

	Earth Matter Effects				Shock effects		ν_e burst vanishing
	ν_e		$\bar{\nu}_e$		ν_e	$\bar{\nu}_e$	
	Peak	Tail	Peak	Tail			
NH, $\sin^2\theta_{13} \gtrsim 10^{-3}$							
Phase A	X	X	√	√	√	X	√
Phase C	X	√	√	X	√	X	√
NH, $\sin^2\theta_{13} \lesssim 10^{-5}$							
Phase A	√	√	√	√	X	X	X
Phase C	√	X	√	X	X	X	X
IH, $\sin^2\theta_{13} \gtrsim 10^{-3}$							
Phase A	X	X	√	√	X	√	X
Phase C	X	√	√	√	X	√	X
IH, $\sin^2\theta_{13} \lesssim 10^{-5}$							
Phase A	X	X	X	X	X	X	X
Phase C	X	√	X	√	X	X	X

Smoking gun signals and caveats

Earth matter effects

- Hierarchy identification even for extremely small θ_{13} values
- If primary fluxes are similar, identifying Earth effects is hard
- Multi-angle effects still to be understood (should not affect in accretion phase)
- Better results with ν_e spectrum $\Rightarrow$ Ar detector crucial

S.Choubey et al., 2010

Shock wave effects

- Presence / absence independent of collective effects
- Stochastic density fluctuations: may partly erase the shock wave imprint G. Fogli et al, 2003
- Turbulent convections behind the shock wave: gradual depolarization effects A. Friedland et al., 2006; J. Kneller et al, 2008, 2010,

Neutronization burst signal

- Robust, but needs Ar detector with good time resolution